

**Supplementary Material for Fletcher et al. A Hexagon in Saturn's Northern Stratosphere Surrounding the Emerging Summertime Polar Vortex**

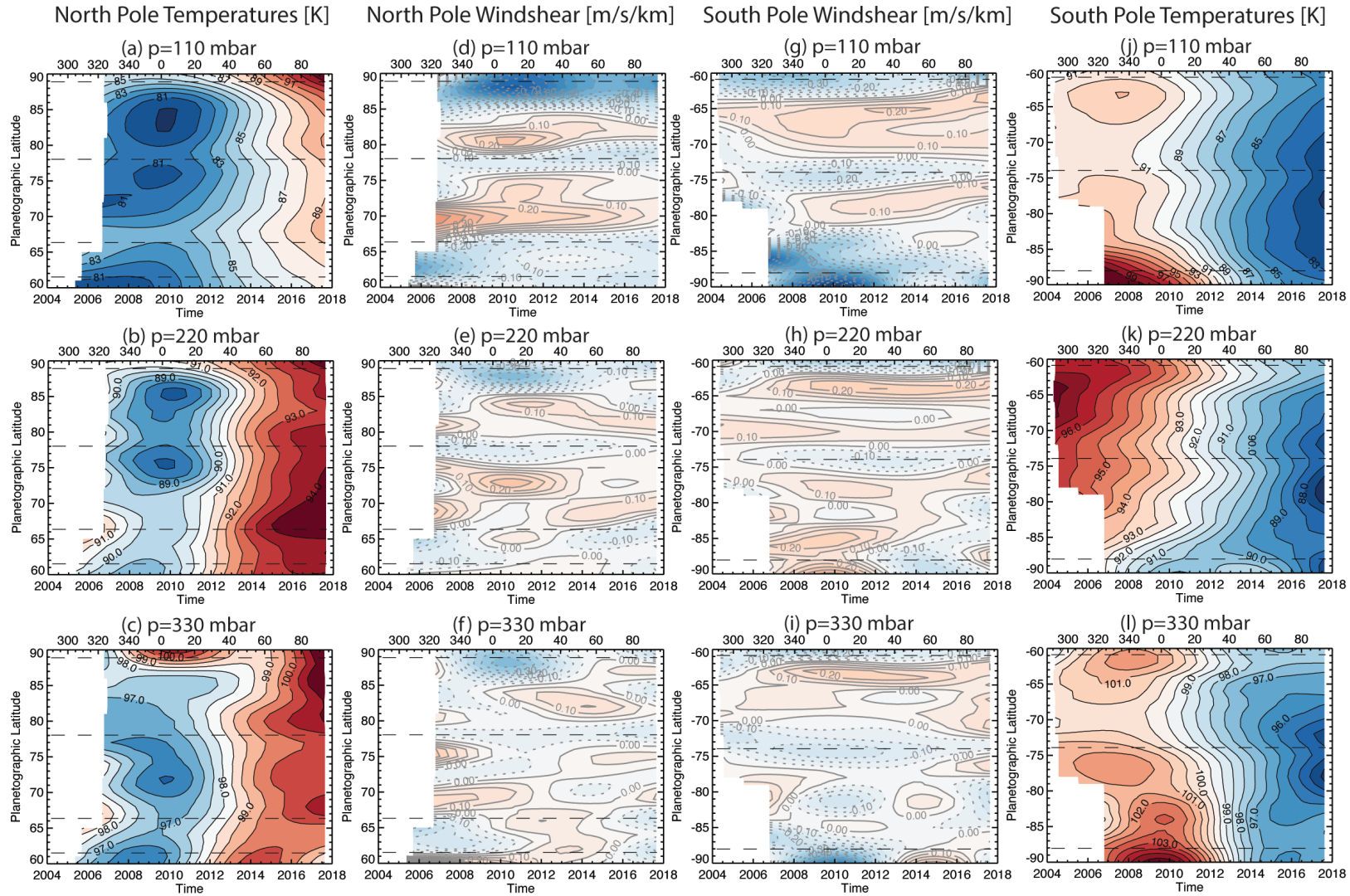


Figure 1: **Supplementary Fig. 1: North and South polar tropospheric temperature gradients as a function of time throughout the whole Cassini mission, 2004-2017.** We display north polar temperatures (panels a-c), north polar windshears (panels d-f), south polar windshears (panels g-i) and south polar temperatures (panels j-l) at three different pressure levels (110, 220 and 330 mbar). These were derived from averages of low-resolution CIRS spectra on a monthly temporal grid, and interpolated using tensioned splines<sup>1</sup> to reconstruct a smoothed temperature field. Horizontal dashed lines signify the peak of eastward zonal jets in the troposphere<sup>2</sup>. The data are displayed as a function of time (years), but a second horizontal axis provides the planetocentric solar longitude ( $L_s$ ) in degrees.

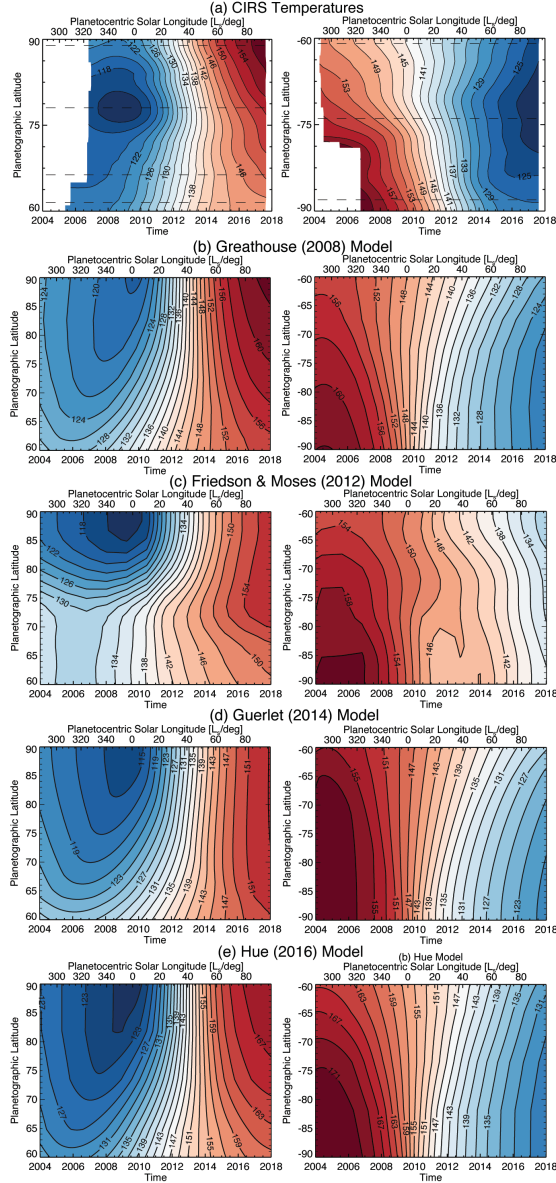


Figure 2: **Supplementary Fig. 2: Comparison of the 1-mbar temperature field to the predictions of radiative models.** We compare (a) the CIRS low-resolution measurements to predictions of (b) the radiative model of Greathouse et al.<sup>3</sup>; (c) the radiative-dynamical model of Friedson and Moses<sup>4</sup>; (d) the radiative-convective model of Guerlet et al.<sup>5</sup>; and (e) the combined radiative-photochemical model of Hue et al.<sup>6</sup>. Results for the north pole are shown on the left, and the south pole are shown on the right. Although the predicted temperatures are close to those observed, none of the models include the warming of the North Pole, and none of them predict the appearance of the strong  $\partial T/\partial y$  gradients associated with the polar vortices.

Table 1: Supplementary Table 1: Sources of spectroscopic linedata and foreign broadening assumptions. Exponents for temperature dependence  $T^n$  given in the final column.

Gas	Line Intensities	Broadening Half Width	Temperature Dependence
CH <sub>4</sub> , CH <sub>3</sub> D	Brown et al. <sup>7</sup>	H <sub>2</sub> broadened using a half-width of 0.059 cm <sup>-1</sup> atm <sup>-1</sup> at 296 K	$n = 0.44$ Margolis et al. <sup>8</sup>
C <sub>2</sub> H <sub>6</sub>	Vander-Auwera et al. <sup>9</sup>	0.11 cm <sup>-1</sup> atm <sup>-1</sup> at 296 K Blass et al. <sup>10</sup>	$n = 0.94$ Halsey et al. <sup>11</sup>
C <sub>2</sub> H <sub>2</sub>	GEISA'03 <sup>12</sup>	Fits to data in Varanasi et al. <sup>13</sup>	Varanasi et al. <sup>13</sup>
PH <sub>3</sub>	Kleiner et al. <sup>14</sup>	Broadened by both H <sub>2</sub> and He using $\gamma_{H_2} = 0.1078 - 0.0014J$ cm <sup>-1</sup> atm <sup>-1</sup> and $\gamma_{He} = 0.0618 - 0.0012J$ cm <sup>-1</sup> atm <sup>-1</sup> Levy et al. <sup>15</sup> , Bouanich et al. <sup>16</sup>	$n = 0.702 - 0.01J$ ( $J$ is the rotational quantum number) Salem et al. <sup>17</sup>
NH <sub>3</sub>	Kleiner et al. <sup>14</sup>	Empirical model of Brown et al. <sup>18</sup>	Brown et al. <sup>18</sup>
C <sub>2</sub> H <sub>4</sub>	GEISA'03 <sup>12</sup>	Polynomial fits to data from Bouanich et al. <sup>19,20</sup>	$n = 0.73$ , Bouanich et al. <sup>20</sup>
C <sub>3</sub> H <sub>4</sub>	GEISA'09 <sup>21</sup>	0.075 cm <sup>-1</sup> atm <sup>-1</sup> for all lines	$n = 0.50$ assumed
C <sub>4</sub> H <sub>2</sub>	GEISA'09 <sup>21</sup>	0.1 cm <sup>-1</sup> atm <sup>-1</sup> for all lines	$n = 0.75$ assumed
C <sub>3</sub> H <sub>8</sub>	GEISA'09 <sup>21</sup>	0.08 cm <sup>-1</sup> atm <sup>-1</sup> for all lines	$n = 0.75$ assumed.
C <sub>6</sub> H <sub>6</sub>	GEISA'09 <sup>21</sup>	Linear fit to N <sub>2</sub> -broadening of Waschull et al. <sup>22</sup>	$n = 0.75$ assumed

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